Compressible Flow - TME085

Chapter 2

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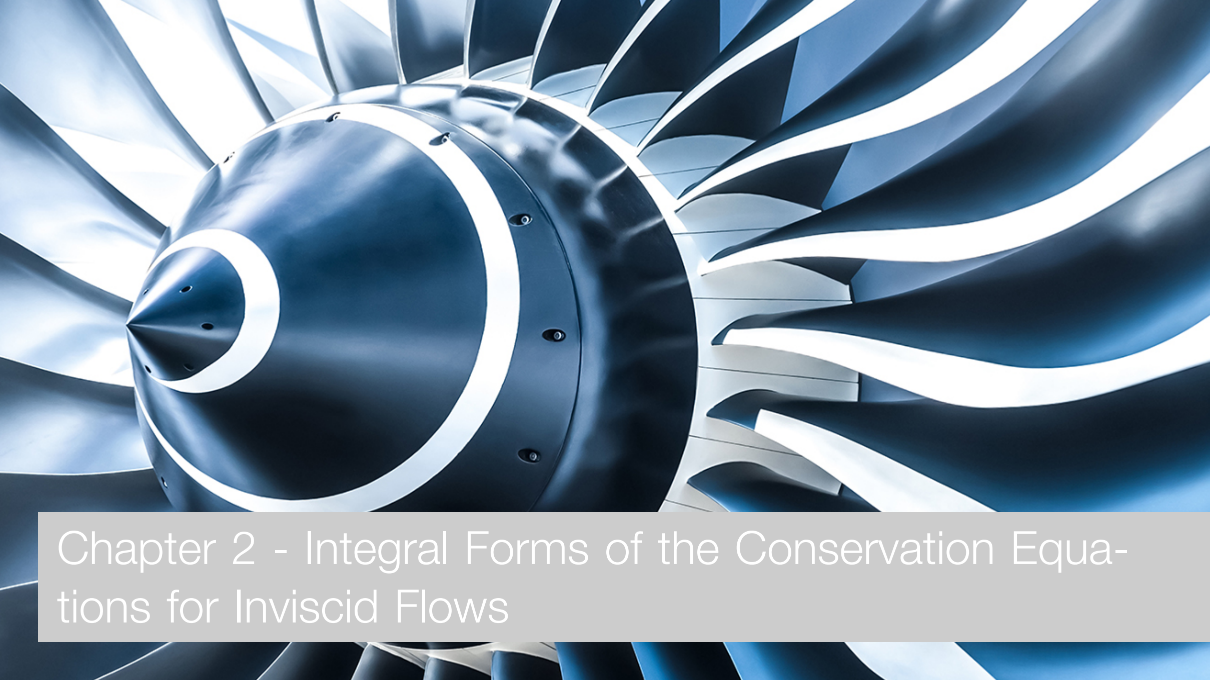
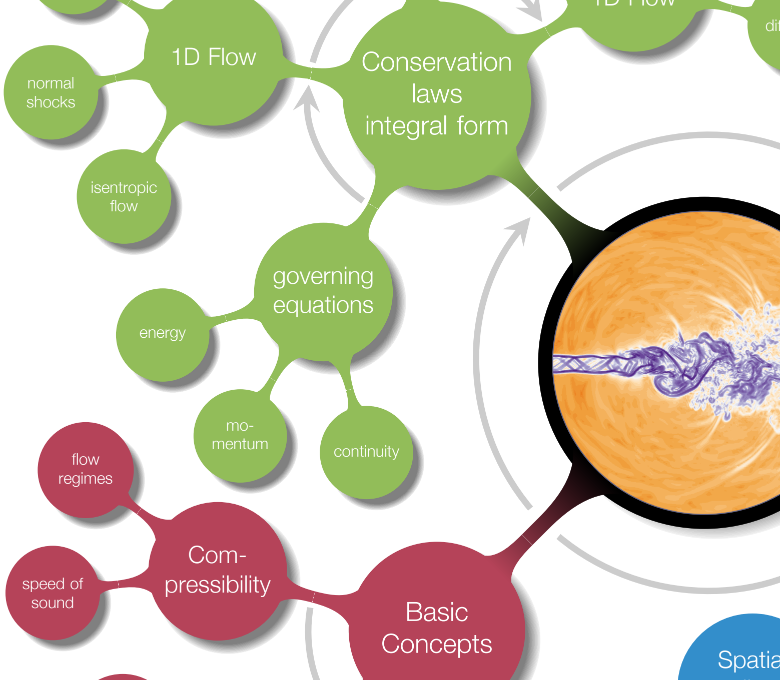


Chapter 2 - Integral Forms of the Conservation Equations for Inviscid Flows

Overview



1D Flow
normal shocks
isentropic flow
Conservation laws integral form
1D Flow
governing equations
energy
mo-mentum
continuity
flow regimes
speed of sound
Com-pressibility
Basic Concepts

Learning Outcomes

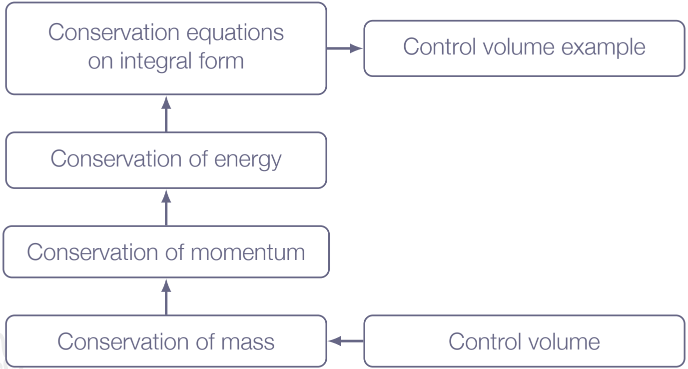
4 **Present** at least two different formulations of the governing equations for compressible flows and **explain** what basic conservation principles they are based on

5 **Explain** how thermodynamic relations enter into the flow equations

7 **Explain** why entropy is important for flow discontinuities

equations, equations and more equations

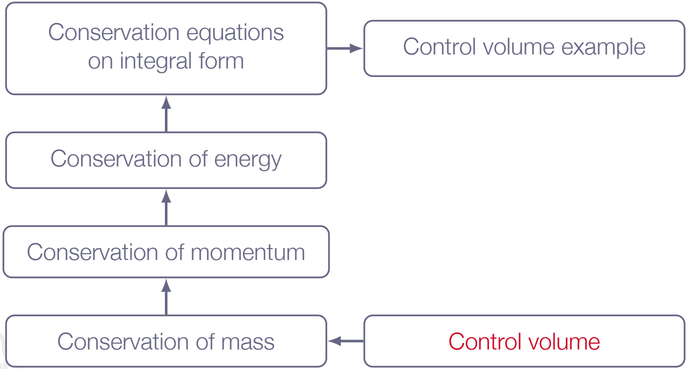
Roadmap - Integral Relations



Motivation

We need to formulate the basic form of the governing equations for compressible flow before we get to the applications

Roadmap - Integral Relations



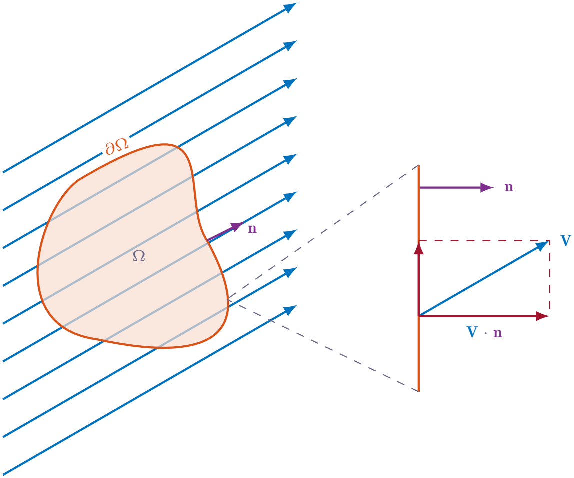
Integral Forms of the Conservation Equations

Conservation principles:

1. conservation of mass
2. conservation of momentum (*Newton's second law*)
3. conservation of energy (*first law of thermodynamics*)

Integral Forms of the Conservation Equations

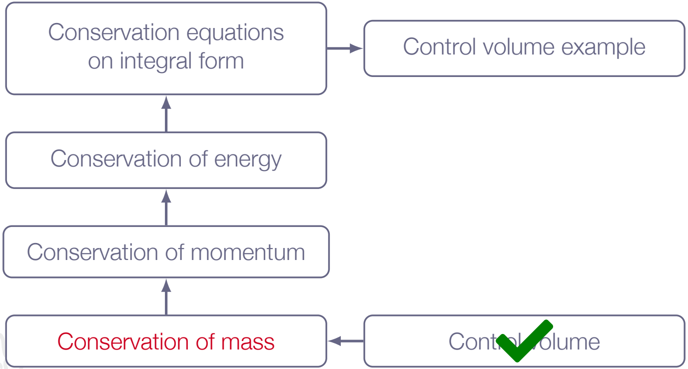
The control volume approach



Notation:

Ω fixed control volume
$\partial\Omega$ boundary of Ω
$\mathbf{n}$ outward facing unit normal vector
$\mathbf{v}$ fluid velocity ($v = |\mathbf{v}|$)

Roadmap - Integral Relations



Chapter 2.3
Continuity Equation

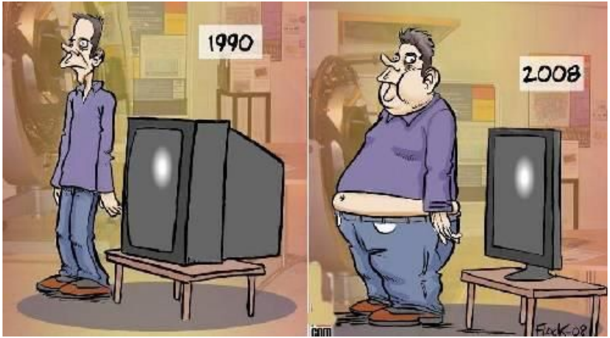
Continuity Equation

Conservation of mass:

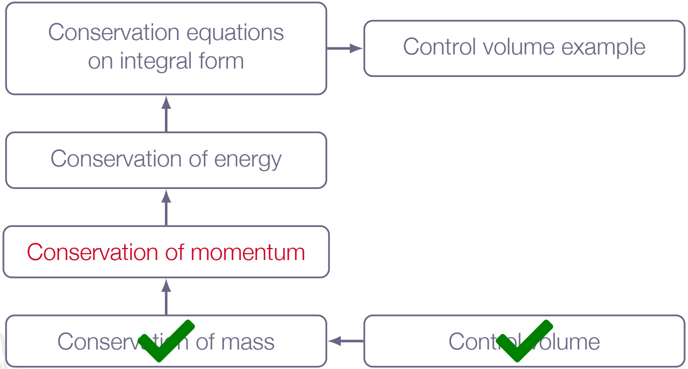
$$\underbrace{\frac{d}{dt}\iiint\limits_{\Omega} \rho d\mathscr{V}}_{\substack{\text{rate of change of}\\ \text{total mass in } \Omega}} + \underbrace{\oiint\limits_{\partial\Omega} \rho \mathbf{v}\cdot\mathbf{n} dS}_{\substack{\text{net mass flow out}\\ \text{from } \Omega}} = 0$$

Note! notation in the text book $\mathbf{n} \cdot dS = d\mathbf{S}$

Conservation of Mass



Roadmap - Integral Relations



Chapter 2.4
Momentum Equation

Momentum Equation

Conservation of momentum:

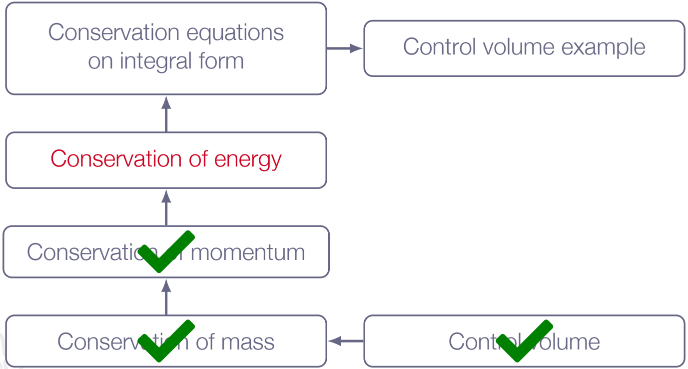
$$\underbrace{\frac{d}{dt}\iiint_{\Omega} \rho \mathbf{v} d\mathscr{V}}_{\substack{\text{rate of change of total} \\ \text{momentum in } \Omega}} + \underbrace{\oiint_{\partial\Omega} \left[\rho(\mathbf{v}\cdot\mathbf{n})\mathbf{v} + p\mathbf{n}\right] dS}_{\substack{\text{net momentum flow out from} \\ \Omega \text{ plus surface force on } \partial\Omega \\ \text{due to pressure}}} = \underbrace{\iiint_{\Omega} \rho \mathbf{f} d\mathscr{V}}_{\substack{\text{rate of momentum} \\ \text{generation due to} \\ \text{forces inside } \Omega}}$$

Note! friction forces due to viscosity are not included here. To account for these forces, the term $-(\tau\cdot\mathbf{n})$ must be added to the surface integral term.
The body force, f, is force per unit mass.

Newton


1. Explain Newton's First Law of Motion in your own words.
!
YAKKA FOOB MOG. GRUG PUBBAWUP ZINK WATTOOM GAZORK. CHUMBLE SPUZZ.
I LOVE LOOPHOLES.

Roadmap - Integral Relations



Chapter 2.5
Energy Equation

Energy Equation

Conservation of energy:

$$\underbrace{\frac{d}{dt}\iiint\limits_{\Omega} \rho e_o d\mathscr{V}}_{\substack{\text{rate of change of total}\\ \text{internal energy in } \Omega}} + \underbrace{\oiint\limits_{\partial\Omega} \left[\rho e_o(\mathbf{v}\cdot\mathbf{n}) + p\mathbf{v}\cdot\mathbf{n}\right] dS}_{\substack{\text{net flow of total internal energy}\\ \text{out from } \Omega \text{ plus work due to}\\ \text{surface pressure on } \partial\Omega}} = \underbrace{\iiint\limits_{\Omega} \rho\mathbf{f}\cdot\mathbf{v} d\mathscr{V}}_{\substack{\text{work due to forces}\\ \text{inside } \Omega}}$$

where

$$\rho e_o = \rho\left(e + \frac{1}{2}\mathbf{v}\cdot\mathbf{v}\right) = \rho\left(e + \frac{1}{2}v^2\right)$$

is the total internal energy

Energy Equation

The surface integral term may be rewritten as follows:

$$\oiint_{\partial\Omega} \left[\rho\left(e + \frac{1}{2}v^2\right)(\mathbf{v}\cdot\mathbf{n}) + p\mathbf{v}\cdot\mathbf{n}\right] dS$$

$$\Leftrightarrow$$

$$\oiint_{\partial\Omega} \left[\rho\left(e + \frac{p}{\rho} + \frac{1}{2}v^2\right)(\mathbf{v}\cdot\mathbf{n})\right] dS$$

$$\Leftrightarrow$$

$$\oiint_{\partial\Omega} \left[\rho\left(h + \frac{1}{2}v^2\right)(\mathbf{v}\cdot\mathbf{n})\right] dS$$

Energy Equation

Introducing total enthalpy

$$h_o = h + \frac{1}{2}v^2$$

we get

$$\frac{d}{dt}\iiint_{\Omega} \rho e_o d\mathscr{V} + \oiint_{\partial\Omega} \left[\rho h_o \mathbf{v}\cdot\mathbf{n}\right] dS = \iiint_{\Omega} \rho\mathbf{f}\cdot\mathbf{v} d\mathscr{V}$$

Energy Equation

Note 1: to include friction work on $\partial\Omega$, the energy equation is extended as

$$\frac{d}{dt}\iiint_{\Omega} \rho e_o d\mathscr{V} + \oiint_{\partial\Omega} [\rho h_o \mathbf{v}\cdot\mathbf{n} - (\tau\cdot\mathbf{n})\cdot\mathbf{v}]\, dS = \iiint_{\Omega} \rho\mathbf{f}\cdot\mathbf{v} d\mathscr{V}$$

Note 2: to include heat transfer on $\partial\Omega$, the energy equation is further extended

$$\frac{d}{dt}\iiint_{\Omega} \rho e_o d\mathscr{V} + \oiint_{\partial\Omega} [\rho h_o \mathbf{v}\cdot\mathbf{n} - (\tau\cdot\mathbf{n})\cdot\mathbf{v} + \mathbf{q}\cdot\mathbf{n}]\, dS = \iiint_{\Omega} \rho\mathbf{f}\cdot\mathbf{v} d\mathscr{V}$$

where $\mathbf{q}$ is the heat flux vector

Energy Equation

Note 3: the force $\mathbf{f}$ inside Ω may be a distributed body force field

Examples:

Gravity

Coriolis and centrifugal acceleration terms in a rotating frame of reference

Energy Equation

Note 4: there may be objects inside Ω which we choose to represent as sources of momentum and energy.

For example, there may be a solid object inside Ω which acts on the fluid with a force $\mathbf{F}$ and performs work $\dot{W}$ on the fluid

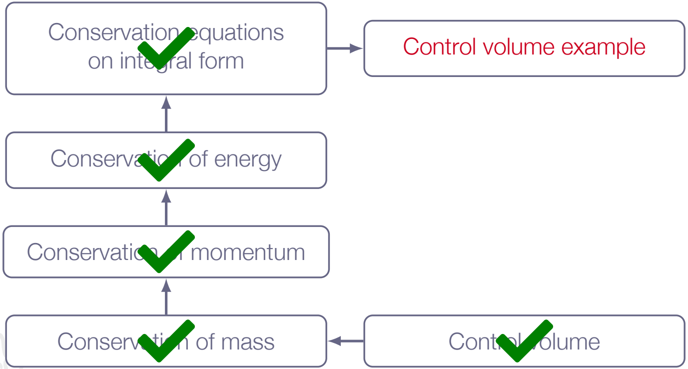
Momentum equation:

$$\frac{d}{dt}\iiint\limits_{\Omega} \rho \mathbf{v} d\mathscr{V} + \oiint\limits_{\partial\Omega} [\rho(\mathbf{v}\cdot\mathbf{n})\mathbf{v} + p\mathbf{n}]\, dS = \iiint\limits_{\Omega} \rho \mathbf{f} d\mathscr{V} + \mathbf{F}$$

Energy equation:

$$\frac{d}{dt}\iiint\limits_{\Omega} \rho e_o d\mathscr{V} + \oiint\limits_{\partial\Omega} [\rho h_o \mathbf{v}\cdot\mathbf{n}]\, dS = \iiint\limits_{\Omega} \rho \mathbf{f}\cdot\mathbf{v} d\mathscr{V} + \dot{W}$$

Roadmap - Integral Relations



Integral Equations - Applications

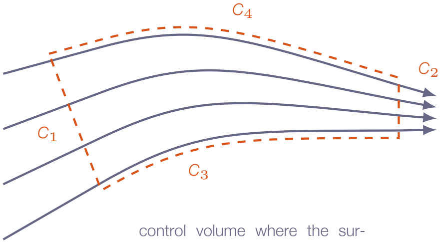
How can we use control volume formulations of conservation laws?

Let $\Omega \to 0$: In the limit of vanishing volume the control volume formulations give the **P**artial **D**ifferential **E**quations (**PDE**:s) for mass, momentum and energy conservation (see Chapter 6)

Apply in a ”smart” way $\Rightarrow$ Analysis tool for many practical problems involving compressible flow (see Chapter 2, Section 2.8)

Integral Equations - Applications

Example: Steady-state adiabatic inviscid flow



control volume where the surfaces C_1 and C_2 are normal to the flow and C_3 and C_4 are parallel to the stream lines

Integral Equations - Applications

Conservation of mass:

$$\underbrace{\frac{d}{dt}\iiint\limits_{\Omega} \rho d\mathscr{V}}_{=0} + \underbrace{\oint\!\!\!\oint\limits_{\partial\Omega} \rho \mathbf{v} \cdot \mathbf{n} dS}_{-\rho_1 v_1 A_1 + \rho_2 v_2 A_2} = 0$$

Conservation of energy:

$$\underbrace{\frac{d}{dt}\iiint\limits_{\Omega} \rho e_o d\mathscr{V}}_{=0} + \underbrace{\oint\!\!\!\oint\limits_{\partial\Omega} \left[\rho h_o \mathbf{v} \cdot \mathbf{n}\right] dS}_{-\rho_1 h_{o_1} v_1 A_1 + \rho_2 h_{o_2} v_2 A_2} = 0$$

Integral Equations - Applications

Conservation of mass:

$$\rho_1 v_1 A_1 = \rho_2 v_2 A_2$$

Conservation of energy:

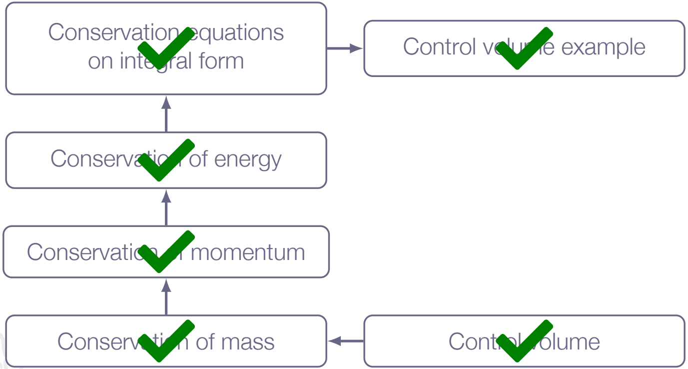
$$\rho_1 h_{o_1} v_1 A_1 = \rho_2 h_{o_2} v_2 A_2$$

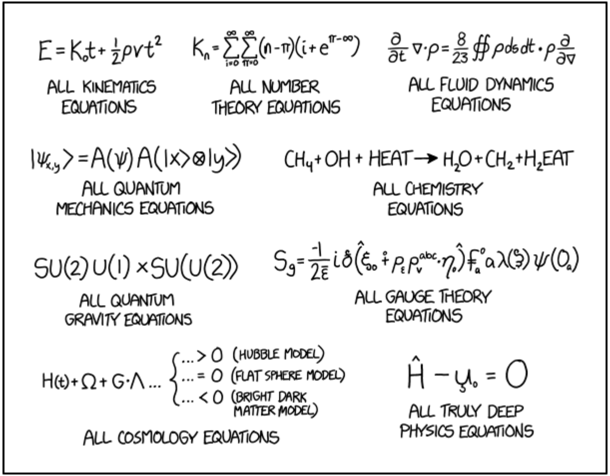
$$\Leftrightarrow$$

$$h_{o_1} = h_{o_2}$$

Total enthalpy h_o is conserved along streamlines in steady-state adiabatic inviscid flow

Roadmap - Integral Relations





$E = K_0 t + \frac{1}{2}\rho v t^2$
ALL KINEMATICS EQUATIONS
ALL NUMBER THEORY EQUATIONS
ALL FLUID DYNAMICS EQUATIONS
$|\psi_{x,y}\rangle = A(\psi)A(|x\rangle \otimes |y\rangle)$
ALL QUANTUM MECHANICS EQUATIONS
$CH_4 + OH + HEAT \rightarrow H_2O + CH_2 + H_2EAT$
ALL CHEMISTRY EQUATIONS
$SU(2)\,U(1) \times SU(U(2))$
ALL QUANTUM GRAVITY EQUATIONS
ALL GAUGE THEORY EQUATIONS
$H(t) + \Omega + G \cdot \Lambda \ldots$
$\ldots > 0$ (HUBBLE MODEL)
$\ldots = 0$ (FLAT SPHERE MODEL)
$\ldots < 0$ (BRIGHT DARK MATTER MODEL)
ALL COSMOLOGY EQUATIONS
$\hat{H} - \mathsf{y}_0 = 0$
ALL TRULY DEEP PHYSICS EQUATIONS